M.G.S. ACCELEROMETER DATA ANALYSIS WITH THE LMD GCM.


Introduction

Mars Global Surveyor aerobreaking phases, required to achieve its mapping orbit, have yielded vertical profiles of thermospheric densities, scale heights and temperatures covering a broad range of local times, seasons and spatial coordinates [Keating et al. 1998, 2001]. Phase I covered local times from 11 to 16 h (assuming 24 "martian hours" per martian day or sols), with a latitude coverage of approximately 40° to 60°N. Seasons observed during this phase were centered around winter solstice and altitudes of periapsis range from 115 to 135 km. The altitudes for Phase II were lower, with a minimum around 100 km and a maximum around 120. Martian spring was the season covered during this phase and the local time was between 15 and 16 h. The latitude covered by Phase II, however, was more extensive than that seen during Phase I, with a coverage from 60°N to basically the South Pole.

MGS density data above 110 km as sampled at a fixed local time yield large amplitude longitudinal variations composed of a variety of wavenumbers (s) which are characterized by large s = 2 component and smaller s = 1 and s = 3 contributions. These oscillations were first interpreted as possible stationary planetary waves [Keating et al. 1998], while it has been later suggested that eastward propagating waves of diurnal frequency could be responsible for the density signature [Forbes and Hagan, 2000; Wilson, 2000].

A fixed local time sampling of the atmosphere will yield a signature composed of a variety of different waves. Determining the dominant ones can be an impossible feat from the observations alone. Waves that propagate with the apparent motion of the sun and have periods that are subharmonics of a day (i.e. the migrating tides) will yield a constant value and no longitudinal variation when sampled at a fixed local time. Only non-migrating tides will contribute to a longitudinal structure at a given local time.

By nonmigrating tides, one refers to those oscillations that have periods that are subharmonics of a day, but their phase speed (\( \zeta \)) differs from the migrating tide’s value of \( \frac{2 \pi}{L} \) (hr\(^{-1} \)). These oscillations, in contrast to the westward migrating tides, can have eastward or westward propagation with respect to the sun, or be standing. Nonmigrating tides can be directly forced, or arise as the outcome of a nonlinear interaction between other preexisting waves.

GCM Simulations and MGS Data Comparisons:

The Mars LMD GCM has evolved from a terrestrial climate model [Hourdin et al., 1993; Forget et al., 1999] and has been extended to the Mars upper atmosphere through a collaboration with the University of Oxford and the Instituto de Astrofisica de Andalucia. The model extends from the ground up to a height of approximately 120 km and includes the relevant processes near the surface such as turbulent diffusion in the planetary boundary layer, convection, orography and low-level drag, all at sub-grid scales. The energetics of the simulations include the effects of the presence in the Mars atmosphere of dust and CO\(_2\). Parameterizations of thermal infrared cooling and near-infrared heating by CO\(_2\) both account for non-LTE effects. The CO\(_2\) condensation–sublimation cycle is also realistically included. A more detailed description of the model can be found in Forget et al. [1999].

MGS Accelerometer observations used in this analysis are NASA Level 3 data set consisting of atmospheric properties at constant altitude levels for each orbital pass covering from 130 to 160 km height [Keating et al., 2001]. Given that lower altitudes were required since our model simulations don’t reach such heights, periapsis density data was scaled to the four following levels: 110, 115, 120 and 130 km. This latter height was only obtained for the orbits that had periapsis altitude above 122 km. Data was sorted according to common local times at each altitude and the profiles obtained, where a significant number of points was available, are shown in Figure 1 along with GCM results that correspond to the observations.

In general, model simulations are able to reproduce the observations with very good agreement. Notable agreement is found at 115 km for the northern spring seasons of Ls=30°-60° (Figure 1a) and Ls=60°-90° (Figure 1e). For these two cases the wave structure observed from orbit is clearly reproduced in the GCM data, and seems to be strongly composed of wavenumbers 2 and 3. Comparisons at Ls=60°-90° reveal a latitudinal dependence, with lesser agreement with the observations in the southern hemisphere as the latitude becomes more polar. This can be seen in Figure 1b to 1d. A plausible explanation for this discrepancy is the large day-to-day variability that is found to exist at high southern latitudes above 90 km in the GCM simulations. The positioning of this highly variable region over the top of the eastward zonal jet is suggestive of a baroclinic instability generating some of the wave activity present there. At 130 km
Figure 1: MGS versus LMD GCM density comparison. MGS data interpolated at the altitude Z (above the grid) are represented with solid dots, while GCM results obtained at the same location and time are displayed in diamonds. Units are kg/km$^3$. Season, time and location are marked for each panel.
(Figure 1f), some of the differences could be ascribed to the high altitude of the observations, this being almost the limit of the valid simulation data. Figures 1g and 1h show comparisons at 120 km during two dusty seasons around northern winter solstice. The agreement with the observations is rather good and slight differences can possibly be due to regional dust storm properties of the period.

### Wave Analysis

Given the rather striking agreement obtained between MGS data and the GCM simulations for the same conditions, an analysis of the modeled data can shed some light on the wave activity at work in the upper atmosphere of Mars. The focus of our analysis is on Phase II MGS accelerometer data. These observations correspond to orbits 582 to 930 which saw a progression of aero-centric longitude from \( L_s = 33^\circ \) to \( L_s = 73^\circ \). For this orbit period, the latitude of periapsis saw a southward progression from 60°N to approximately 20°S. The local solar time sampled a two-hour bracket from 17 to 15 h. Specifically, Figure 1e represents equatorial latitudes (10° S-22° N) at a height of 115 km for \( L_s \approx 70^\circ \), which correspond to a local time of 15 h. As can be seen, a very good agreement with the observations is obtained with the wave structure clearly represented. The data shows a possible combination of wavenumber 2 and 3 as seen in a fixed local time frame of reference.

Such kind of structures have been observed and interpreted as stationary waves [Keating et al., 1998] and have later been shown to be nonmigrating tides, and more specifically, diurnal Kelvin Waves [Forbes and Hagan, 2000; Wilson, 2002]. The excitation mechanism for these eastward waves has been suggested to lie in the nonlinear modulation of the migrating diurnal tide forcing by Mars’ topography components [Wilson and Hamilton, 1996; Forbes and Hagan, 2000]. A wave analysis of the LMD GCM simulation data yields other significant contributors to the structures observed by MGS.

A two-dimensional Fourier decomposition was performed on the GCM data to extract the prominent waves present. The analysis was performed on ten days of simulation with wavenumbers ranging from -10 to 10 (westward and eastward, respectively) and periods from 8 to 72 hours. Stationary waves with wavenumbers 1 to 6 were also examined. The obtained density amplitudes for the different wavenumbers normalized by the zonal mean density are depicted in Figure 2 corresponding to diurnal frequencies. Magnitudes are largest for the migrating tide, with significant amplitudes obtained for the diurnal eastward propagating waves with \( s = 1, 2 \) and 3. Lesser amplitudes are found for shorter and zonally symmetric waves. Semidiurnal nonmigrating tides (not shown) have in general smaller amplitudes than those of diurnal frequency, with the \( s = 1 \) and zonally symmetric \( (s = 0) \) waves found to maximize above 80 km at high northern latitudes. Significant magnitudes are also seen at wavenumbers two, three and five for the eastward propagating semidiurnal waves.

The largest amplitudes for nonmigrating tides are obtained for eastward propagating diurnal waves with wavenumbers one, two and three (with magnitudes of 1.36, 0.99, and 1.17 respectively). Their contribution to the agreement is significantly improved with the addition of stationary wave one, two and three (not shown) does not increase significantly the closeness of the decomposition. The spike observed at 230°E longitude is nicely reproduced on the simulated data (Figure 1e) and the addition of stationary wave one, two and three (not shown) does not increase significantly the closeness of the decomposition. The spike observed at 230°E longitude is nicely reproduced on the simulated data (Figure 1e) and the agreement is significantly improved with the addition of these three waves.
crepancies are thought to relate to differences in the dust distribution in the simulations. Local time for Ls=70 the main features observed by MGS, although few dis-
latitude comparisons present larger differences between
the modeled and observed data. These differences are
served data which allows for its use in its analysis. Results
braking accelerometer data from MGS for various sea-
sulations fare very well in the comparison with the ob-
observations since a variety of wavenumber-frequency
combinations can yield a certain signature when sampled
at a constant local time. This makes the use of model
simulations an invaluable tool in the understanding of
observations.

Two-dimensional Fourier decomposition of the sim-
ulation data yields the diurnal eastward propagating
waves with wavenumber one, two and three as the ma-
jor contributors to the structure observed by MGS at equatorial latitudes. Other significant waves obtained are the semidiurnal eastward propagating waves with wavenumbers two and five which along with the diur-
nal s = 5 help create the strong signature at a longitude of 230°E seen in the data. It is striking that such a combination of minor waves can increase the density encountered by a spacecraft by almost 40%. This illustrates the complex nature of the Martian atmosphere at those heights and the need for further observations and continuing modelling efforts.

Acknowledgements

This research was funded under cNES Mars Premier Program and ESTEC Contract 11369/95/NLJ.

References


